

On the stability of the thermal Comptonization index in neutron star low-mass X-ray binaries in their different spectral states

R. Farinelli¹ & L. Titarchuk^{1,2}

¹ Dipartimento di Fisica, Università di Ferrara, Via Saragat 1, 44100 Ferrara, Italy

² NASA/GSFC, Greenbelt, MD 20771

Preprint online version: September 7, 2010

ABSTRACT

Context. Most of the spectra of neutron star low mass X-ray binaries (NS LMXBs), being them persistent or transient, are characterized by the presence of a strong thermal Comptonization bump, thought to originate in the transition layer (TL) between the accretion disk and the NS surface. The observable quantities which characterize this component dominating the emission below 30 keV, are the spectral index α and the rollover energy, both related to the electron temperature and optical depth of the plasma.

Aims. Starting from observational results on a sample of NS LMXBs in different spectral states, we formulate the problem of X-ray spectral formation in the TL of these sources. We predict a stability of the thermal Comptonization spectral index in different spectral states if the energy release in the TL is much higher than the intercepted flux coming from the accretion disk.

Methods. We use an equation for the energy balance and the radiative transfer diffusion equation for a slab geometry in the TL, to derive a formula for the thermal Comptonization index α . We show that in this approximation the TL electron temperature kT_e and optical depth τ_0 can be written as a function of the energy flux from the disk intercepted by the corona (TL) and that in the corona itself, $Q_{\text{disk}}/Q_{\text{cor}}$. As spectral index α depends on kT_e and τ_0 , this in turn leads to a relation $\alpha = f(Q_{\text{disk}}/Q_{\text{cor}})$, with $\alpha \sim 1$ when $Q_{\text{disk}}/Q_{\text{cor}} \ll 1$.

Results. We show that the observed spectral index α for the sample of sources here considered lies in a belt around 1 ± 0.2 a part for the case of GX 354–0. Comparing our theoretical predictions with observations, we claim that this result, which is consistent with the condition $Q_{\text{disk}}/Q_{\text{cor}} \ll 1$, can give us constraints on the accretion geometry of these systems, an issue that seems difficult to be solved using only the spectral analysis method.

Key words. Stars: neutron — X-rays: binaries — Accretion, accretion disks — Radiative transfer

1. Introduction

It is well known since the 80's that the spectra of low-mass X-ray binaries (LMXBs) hosting a neutron star (NS) can be described up to about 30 keV by a two-component model representing the contribution from different emitting regions of the system. However, the interpretation of the spectra is not unique as demonstrated using the variety of different models used across the years.

Before the *BeppoSAX* and *RXTE* era, two concurring models were classically used to describe X-ray emission in LMXBs. In the so-called “eastern model” (Mitsuda et al. 1984; Mitsuda et al. 1989) the spectra were fitted by the sum of a soft blackbody (BB) emission (actually modeled by a multi-colour disk BB spectrum) attributed to the accretion disk, plus a hotter simple or Comptonized BB claimed to originate close to the NS surface. On the other hand, in the “western model” inter-

pretation (White et al. 1986; White et al. 1988), the direct BB component was attributed to the NS surface, while an unsaturated Comptonization spectrum was thought to originate from a hot corona above the inner accretion disk, which supplies most of the soft seed photons for Comptonization.

In fact, even after the advent of *BeppoSAX* and *RXTE*, the persistent emission of NS LMXBs was described by the sum of a BB component plus a thermal Comptonization (TC) spectrum, usually described by the XSPEC COMPTT model (Titarchuk 1994; Hua & Titarchuk 1995). Despite the significant improvement in our knowledge of the source spectral properties by means of the broad-band observations, the BB+TC model resulted and subjected to a dichotomy. Indeed, both of the cases where the temperature of the direct BB spectrum kT_{bb} is lower (e.g., Di Salvo et al. 2000a,b, 2001; Oosterbroek et al. 2001; Gierliński & Done 2002; Lavagetto et al. 2004; Paizis et al. 2005) or higher (e.g., Paizis et al. 2005; Farinelli et al. 2007; Farinelli et al. 2008,

hereafter F08) than that of the thermally Comptonized seed photons, kT_s , provide in general equally acceptable good fits.

The consequences of these results were the interpretation of the BB emission as due either to the accretion disk ($kT_{bb} < kT_s$) or to the NS surface ($kT_{bb} > kT_s$). Moreover, in addition to the persistent X-ray emission, *BeppoSAX* (e.g., Di Salvo et al. 2000b; Di Salvo et al. 2002) *RXTE* (D’Amico et al. 2001; Di Salvo et al. 2006) and later also *INTEGRAL* (Paizis et al. 2006, hereafter P06) allowed the possibility to discover in bright NS LMXBs a transient power-law (PL) X-ray emission above 30 keV.

Motivated by the need to put some order and give an unified scenario in the different NS systems spectral states, Paizis et al. (2006, hereafter P06) performed a systematic observational campaign with the *ISGRI* (20-200 keV) monitor onboard *INTEGRAL*. Using long-term average spectra and including former *INTEGRAL* results on GX 354–0 (Falanga et al. 2006), P06 classified NS LMXBs into four main states: the *hard/PL*, *low/hard*, *intermediate* and *soft*. The high-energy (> 20 keV) spectra in different sources were interpreted by P06 as the result of the interplay between thermal and bulk Comptonization processes whose relative efficiency is ultimately dictated by the mass accretion rate. At high energies (where the direct BB component is negligible), the *hard/PL* state spectra can be fitted with a simple PL component; the *low/hard* spectra by a TC spectrum of soft ($\lesssim 1$ keV) BB-like photons off an electron population with $kT_e \sim 20 - 30$ keV and $\tau_0 \lesssim 3$; the *intermediate* state spectra by TC spectrum with $kT_e \sim 3 - 5$ keV and $\tau_0 \gtrsim 5$ plus a PL component with photon index $\Gamma \sim 2 - 3$; the *soft* state spectra by a TC component similar to the *intermediate* state but without the high-energy X-ray tail.

From the observational point of view, the quantities directly measurable in the data are the cut-off energy E_c of the dominating TC bump and the spectral slope, parametrized through the energy index α ($=\Gamma - 1$). For pure TC spectra, α is tightly correlated with the plasma temperature kT_e and optical depth τ_0 (Sunyaev & Titarchuk 1980; Titarchuk & Lyubarskij 1995, hereafter TL95), while in the presence of a converging flow (bulk motion), the shape of the velocity field also determines the emerging spectral slope (Titarchuk et al. 1997; Laurent & Titarchuk 1999, F08). In mathematical terms, α represents the index of the system Green’s function (GF), namely as it responds to monochromatic line injection. The resulting emerging spectrum is then obtained as a convolution of the GF with the input seed photon spectrum. We also emphasize that the *hard/PL*, *hard* and *soft* states spectra of NS LMXBs are characterized by the presence of just *one* Comptonization index (see Fig. 4 in P06). In the first case, its value is interpreted as a result of a mixed TC plus BC process, while in the latter two cases, α is derived from pure TC. On the other hand, the *intermediate* state spectra show *two* Comptonization indexes, one is related to the persistent TC component and the other one characterizes the transient PL-like hard X-ray emission.

In Section 2 we give an overview of the theoretical and observational issues related to spectral evolution in X-ray binary systems hosting a NS or a black hole (BH) outlining the differences among the two classes of sources. We subsequently

report on results related to the observed TC index α for a sample of NS sources. In Section 3 we propose a theoretical model based on diffusion formalism for radiative transfer in order to explain the observational results. In Section 4 we discuss the comparison between theory and data, while in Section 5 we draw our conclusions and give future observational prospects.

2. Spectral index evolution in NS systems

Starting from the considerations of the previous section, it is important to make a comparison between the index evolution in accreting BH and NS sources. The spectral state of BH sources may be generally divided into *low/hard* state (LHS), where the spectrum is dominated by a TC component with electron temperature $kT_e \sim 60-100$ keV, *intermediate* state (IS), with a BB bump (presumably coming from the accretion disk) and a superposed PL high-energy component, and *high/soft* state (HSS) where the BB component gets even stronger and the PL emission gets steeper.

Actually, in BH systems generally *one high energy photon index* Γ is observed [see e.g. recent results on GRS 1915+015 by Titarchuk & Seifina (2009) and Shaposhnikov & Titarchuk (2009)] and its value evolves with the source mass accretion rate. The latter, which is not a direct observable quantity, may be inferred in an indirect way by means of the the normalization of the disk flux N_{disk} . In a Γ vs N_{disk} /Quasi Periodic Oscillation (QPO) frequency diagram, it was found (Shaposhnikov & Titarchuk 2009; Titarchuk & Seifina 2009; Montanari et al. 2009) that the photon index Γ progressively increases from $\sim 1.6-1.8$ as the source moves from the *low/hard* to the *high/soft* states, until it reaches saturation around $\Gamma \sim 2.2 - 2.4$ (depending on the source). This interpretation of the observed BH spectra have been performed using “so called” the BMC model (Titarchuk et al. 1997), whose emerging spectral shape is described by the formula

$$F(E) = \frac{C_N}{A+1} [BB(E) + A \times G(E, E_0) * BB(E_0)]. \quad (1)$$

In formula (1), the first term of the right-hand side represents the direct seed photon BB-like spectrum, while the second one gives its modification due to Comptonization (convolution with GF). The hardness of the spectrum is dictated by the energy index α of the GF which, in BMC, is a broken-PL with $G(E, E_0) \propto E^{\alpha+3}$ for $E < E_0$ and $G(E, E_0) \propto E^{-\alpha}$ for $E > E_0$, where E_0 is the monochromatic input energy. It is important to keep in mind that the GF in BMC does not contain the exponential spectral rollover, so when the latter is observed in the X-ray spectrum of a source, it can be taken into account by multiplying BMC with an exponential e-folding factor $\propto e^{-E/E_c}$.

The index saturation observed in BH sources as they move to the HSS finds a natural explanation in the framework of a bulk-dominated scenario (see mathematical proof of this statement in e.g. Titarchuk & Seifina 2009). Moreover, it is interesting to note that when the sources are in the LHS, the points in the Γ vs N_{disk} /QPO diagram form a plateau around $\Gamma \sim 1.5$ before the rising phase. This clear mapping of the energy (or photon) index evolution is possible in BH sources as they are

generally strongly variable and their X-ray spectrum allows a one-to-one correspondence between the spectral state and the energy index α (or Γ) using spectral modeling according to formula (1). One of the advantages of using BMC model is that it is a *generic* Comptonization model, namely it allows to map the spectral evolution of sources through the Comptonization index α , which is a direct measurable quantity, no matter which are the underlying physical conditions. The theoretical interpretation of the source spectral formation is in fact related to a later scientific discussion.

For NS sources however, situation is less straightforward to face. In fact, most variable sources such as, e.g., GX 354–0 (Falanga et al. 2006) or 4U 1608–52 (Gierliński & Done 2002) when moving from the *hard* state to the *soft* state do exhibit pure TC spectra, with electron temperature kT_e progressively decreasing and optical depth τ_0 increasing, respectively. On the other hand, bright LMXBs of the GX class such as the classical six known Z sources (Sco X–1, GX 17+2, Cyg X–2, GX 340+0, GX 5–1 and GX 349+2) and, more recently, GX 13+1, show only a small evolution of their persistent X-ray continuum (dominated by the strong TC bump with $kT_e \sim 3$ –5 keV and $\tau_0 \gg 1$). In addition they are characterized by the presence of a transient hard X-ray PL-like component (*intermediate* state).

Other persistently bright sources such as GX 3+1, GX 9+1 and GX 9+1 have been only observed in the *soft* state with no evidence of hard X-ray tails. In fact, mapping the evolution of the transient hard X-ray tail of NS sources in the *intermediate* state has not been yet possible because of the insufficient statistics available at high energies. Some attempts (Di Salvo et al. 2000a; Di Salvo et al. 2002, 2006) were actually undertaken to establish changes in the intensity of the transient hard tail by fitting it with a simple PL and, when statistics was poor, fixing the index allowing to vary only its flux. But nothing could be concluded about the index evolution. Thus, at the present status of knowledge, serious investigations can be performed only on the evolution of the spectral index related to the persistent TC component. This was done for the first time by Titarchuk & Shaposhnikov (2005, hereafter TS05), who performed a systematic analysis of the variable NS X-ray binary 4U 1728–34 from the *hard* state to the *soft* state using observations from the Proportional Counter Array (PCA, 3–30 keV) onboard *RXTE*. The source spectra were fitted with a two-BMC model, with $A \gg 1$ (see formula [1]) and with the GF spectral index fixed equal for both BMC components. Thus, the model used by TS05 actually was

$$F(E) = C_{ns}G(E, E_0)*BB(E_0)+C_{disk}G(E, E_0)*BB(E_0).$$

The main result found by TS05 was that as the source moved from the *hard* to the *soft* state the index Γ ($= \alpha+1$) progressively increased with no evidence of saturation, unlike the BH case. Actually, the final *soft* state of 4U 1728–34 was represented by the sum of two BB component, because for $\alpha \gg 1$, $G(E, E_0) * BB(E_0) \approx BB(E)$.

This result however needs a revision. Namely, fitting the NS LMXBs *soft* state spectrum with a two-BB model can be possible because of the lack of data below 3 keV, which is of key

importance. The limited broad-band resolution allows however the possibility to have different models with same good fitting results. For example, recent results of the analysis of PCA/*RXTE* data for 4U 1608–52 by Ding et al. (2010 in preparation) show that the *soft* state spectrum of the source can be fitted either by a two-BB model or by a single TC model (e.g., COMPTB) with $\alpha \sim 1$. Note also that Falanga et al. (2006) fitted the *soft state* spectrum of GX 354–0 with a DBB+COMPTT model, from which the inferred α -value is again ~ 1 . Actually, when looking at the *BeppoSAX* results obtained over years of observations of NS LMXB sources, it results that the *soft* state spectra of these systems, rather than 2 BBs, need to be described by the sum of a BB component plus an unsaturated TC spectrum with cut-off energy below 10 keV.

2.1. Observational results

We considered a sample of sources taken from the literature and for which we can make a fiducial measurement of α . In our choice of the sources, we adopted the criterion to consider those in which α was determined either directly from the fit, as it can be done using the COMPTB model (where α is a free parameter, see F08), or can be derived from the temperature and optical depth obtained by the COMPTT model (Titarchuk 1994).

The sources belonging to the first case are Sco X–1, GX 17+2, Cyg X–2, GX 340+0, GX 3+1 (see Table 2 in F08) and GS 1826–238 (see Table 2 in Cocchi et al. 2010). For GX 349+2 (Di Salvo et al. 2000a) and GX 354–0 (Di Salvo et al. 2000b) we derived the value of spectral index α using equation for the non-relativistic regime (see Eq. [22] in Titarchuk & Lyubarskij 1995, hereafter TL95)

$$\alpha = -\frac{3}{2} + \sqrt{\frac{9}{4} + \frac{\beta}{\Theta}}, \quad (2)$$

where $\Theta \equiv kT_e/m_e c^2$ and β -parameter defined in formula (17) of TL95 for spherical geometry as it was assumed by the authors. In the case of X1658–298, Oosterbroek et al. (2001) assumed a slab geometry, thus α was obtained from equation (2) and formula (17) of TL95 for a slab geometry. For 1E 1724–3045, Barret et al. (2000) report the best-fit value of optical depth τ_0 of Comptonization region both for the case of spherical and slab geometry. We checked that the two derived values of α are perfectly consistent. The errors on α for sources for which kT_e and τ_0 were reported, have been computed considering that the function $\alpha[kT_e, \beta(\tau_0)]$ gets its absolute minimum and maximum values at the boundary of the box of its domain delimited by the minimum and maximum value (kT_e^{\min} , kT_e^{\max}) and (τ_0^{\min} , τ_0^{\max}) obtained in computing the errors at 90% confidence level for the electron temperature and optical depth by XSPEC.

In Figure 1 we report the measured values of α for this sample of sources as a function of the electron temperature kT_e . This parameter can be considered indeed a good tracer of the source spectral state because kT_e decreases when sources move from the *hard* to *soft* state as a result of a more efficient electron cooling by the enhanced seed photon supply. Moreover,

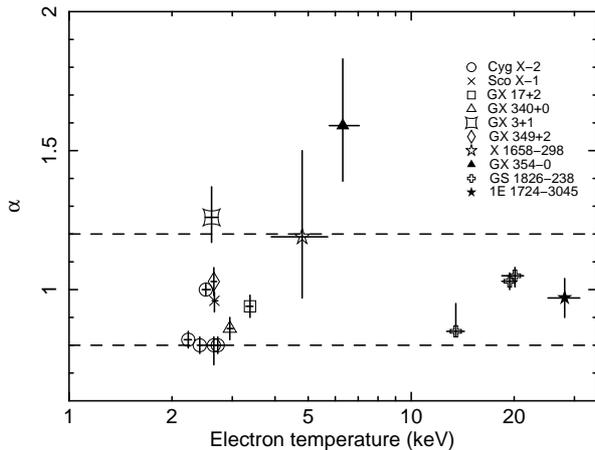


Fig. 1. Thermal Comptonization index α for sources in different spectral states as a function of the electron temperature kT_e . Reference papers: Cyg X-2, Farinelli et al. (2009); Sco X-1, GX 17+2, GX 340+0 and GX 3+1, Farinelli et al. (2008); GX 354-0, Di Salvo et al. (2000b); GX 349+2, Di Salvo et al. (2001), X 1658-298, Oosterbroek et al. (2001); GS 1826-238, Cocchi et al. (2010); 1E 1724-3045, Barret et al. (2000).

the electron temperature kT_e is a directly measurable quantity because it is related to the cut-off energy of the spectrum and it presents the advantage of being distance-independent. On the other hand, the instrumental energy-band coverage and accumulation time can play some role in biasing the measured index value. For instance, Falanga et al. (2006) performed a systematic analysis of GX 354-0 as a function of its position on the hardness-intensity diagram when the source moved from the *hard* to *soft* state. The authors fitted the 3-100 keV spectrum with a multicolour disk BB (DBB, Mitsuda et al. 1984) plus COMPTT model using the slab geometry. We computed the derived α -values from their best-fit parameters but the trend was not monotonic, covering the range ~ 1 -3 and reflecting such a behaviour of the $kT_e - \tau_0$ parameters (decreasing of kT_e was not followed by increasing of τ_0 , as expected). It is not clear whether this was due to the lack of data below 3 keV, which is necessary to constrain the seed photon temperature or due to accumulation time. Thus we prefer to skip these measurements. It is also worth mentioning that Oosterbroek et al. (2001) performed a *BeppoSAX* analysis of GX 3+1 and Ser X-1, both sources characterized by typical *soft* state spectra, fitting them with a DBB+COMPTT model, but they did not specify which geometry (sphere or slab) it was assumed. We found $\alpha_{sph} = 2.44^{+0.89}_{-0.59}$ and $\alpha_{sl} = 0.87^{+0.42}_{-0.25}$ for GX 3+1, $\alpha_{sph} = 1.41^{+0.52}_{-0.34}$ and $\alpha_{sl} = 0.45^{+0.20}_{-0.12}$ for Ser X-1, respectively.

Looking at Fig. 1 we note that for all the analyzed sources the spectral index α lies in a belt around 1 ± 0.2 , a part for the case of GX 354-0 where $\alpha \sim 1.6$. We give a possible interpretation of these *observational* results in the Discussion.

3. A model of Comptonization region in the neutron star case

The determination of spectral index α obtained from Comptonization of seed photons in a bounded medium has been faced since a long time. The emerging radiation spectrum depends on several parameters such as the geometry of the plasma (e.g., slab or sphere), the electron temperature and optical depth, and the space distribution of the seed photons inside the medium. Sunyaev & Titarchuk (1980) report the value α obtained from the solution of the stationary radiative transfer equation in the non-relativistic case (Fokker-Planck approximation) obtained as a convolution of the time-dependent equation with the time-escape probability distribution $P(u)$ for the case of source photons distributed according to the first eigenfunction of the space operator. Later Titarchuk (1994), Hua & Titarchuk (1995) and TL95 extended the results to the sub-relativistic case considering both slab and spherical geometry.

In order to understand what it does happen in NS LMXBs sources, one has to consider the hydrodynamical conditions in the region between the accretion disk and the NS surface. We will refer to this region as the transition layer (TL), often also called boundary layer. Actually, the production of a strong TC bump in the persistent X-ray spectra of NS LMXBs is thought to originate in this TL, namely the region where matter deviates from its Keplerian angular velocity in order to match that of the slowly spinning NS. The radiative and hydrodynamical configuration of the TL is mostly dictated by the Reynolds number γ (Titarchuk et al. 1998; Titarchuk & Osherovich 1999), which is proportional to the mass accretion rate and is eventually the inverse of the viscosity parameter of the Shakura-Sunyaev disk. In particular, γ determines the radial extension of the TL and in turn, from the mass accretion rate, its optical depth. It is worth pointing out that determination of the vertical height-scale of the TL is in fact a very complicated problem, as it should require a complete 3D magneto-hydrodynamical treatment of the problem. Using the slim disk (thus vertical-averaged) equations for determining the radial thermo-hydrodynamical structure of the TL may be in fact an issue (e.g., Popham & Sunyaev 2001).

The enhanced radiation and thermal pressure because of higher electron temperature are expected to increase the vertical height-scale of the TL with $H \approx R_{ns}$. Moreover, the solution for the angular momentum equation (Titarchuk et al. 1998; Titarchuk & Osherovich 1999) for Reynolds number $\gamma > 5 - 10$ gives a TL radial extension $\Delta R_{TL} \lesssim 0.5 R_{NS}$. With these characteristic length-scales, it seems more plausible to approximate the TL geometry to a slab whose normal is directed along the disk plane. It is worth emphasizing that Haardt & Maraschi (1993, hereafter HM93) determined the theoretical Comptonization index α considering a two-phase model for accretion disks in AGN in which a hot corona is surrounding and sandwiching the underlying cold accretion disk. The model can in principle be applied also to the case of solar-mass BH sources. HM93 assume that the corona and the disk are two slabs at significant different temperatures and put in contact each other. The authors concentrated on the case of high tem-

perature ($kT_e \gtrsim 50$ keV) and low optical depth ($\tau < 1$) for the corona, so that the diffusion approximation cannot hold, unlike what we are considering. One of the consequences of the high-temperature treatment is that electron scattering is anisotropic with a significant fraction of the power back-irradiating the disk. In the HM93 model, the *inner* boundary condition of the hot corona is the disk cool surface (with $kT_{bb} < 5$ eV) with energy-dependent albedo. Note also that in their geometry, 100% of the disk flux is intercepted and reprocessed by the top plasma. In the geometry considered here on the other hand, it is possible that part of the disk emission directly escapes the system, while a fraction of its flux is intercepted by the TL. We are actually interested here on this portion of the intercepted disk flux (see next Section).

Because of these differences between the two models, a direct comparison of the derived theoretical results is not straightforward. The reader can refer to the paper of HM93 for further details in order to better understand the differences of our and their approaches.

3.1. Energy release in the NS transition layer

The energy balance in the TL is dictated by Coulomb collisions with protons (gravitational energy release), while inverse Compton and free-free emission are the main cooling channels (see a formulation of this problem in the pioneer work by Zel'dovich & Shakura 1969). In fact, for the characteristics electron temperature (3 keV $\lesssim kT_e \lesssim 30$ keV) and density values ($\lesssim 10^{-5}$ g cm $^{-3}$) of these regions for LMXBs, Compton cooling dominates over free-free emission and the relation between the energy flux per unit surface area of corona Q_{cor} , the radiation energy density $\varepsilon(\tau)$ and electron temperature T_e is given by (see also TLM98)

$$\frac{Q_{cor}}{\tau_0} \approx 20.2\varepsilon(\tau)T_e(\tau), \quad (3)$$

where τ_0 is the characteristic optical depth of the TL.

The distribution $\varepsilon(\tau)$ is obtained as a solution of the diffusion equation

$$\frac{d^2\varepsilon}{d\tau^2} = -\frac{3Q_{tot}}{c\tau_0}, \quad (4)$$

where now $Q_{tot} = Q_{cor} + Q_{disk}$ is the sum of the corona (TL) and intercepted disk fluxes, respectively. The two boundary conditions for equation (4) are written as

$$\left. \frac{d\varepsilon}{d\tau} \right|_{\tau=\tau_0} = 0, \quad (5)$$

$$\left. \frac{d\varepsilon}{d\tau} - \frac{3}{2}\varepsilon \right|_{\tau=0} = 0 \quad (6)$$

which represent the case of albedo $A=1$ at the NS source ($\tau = \tau_0$) and no diffusion emission falling from outside onto outer corona boundary ($\tau = 0$). The condition for $A=1$ arises from the well-established observational result of NS temperature $kT_{bb} \sim 1$ keV, which implies the presence of a ionized NS atmosphere. This is different from the case considered by HM93, where the cool disk temperature (< 5 eV) gives rise to an energy-dependent albedo with photoelectric absorption for

impinging photons with energy $\lesssim 10$ keV. Another important consideration to keep in mind is that the diffusion equation (4) is to be considered frequency-integrated. This means that we are not dealing with the specific (energy-dependent) shape of the reflected spectrum from the NS surface, we are only dealing with the total energy density.

The solution for $\varepsilon(\tau)$ is then given by

$$\varepsilon(\tau) = \frac{2Q_{tot}}{c} \left[1 + \frac{3}{2}\tau_0 \left(\frac{\tau}{\tau_0} - \frac{\tau^2}{2\tau_0^2} \right) \right]. \quad (7)$$

It is worth pointing out that $d\varepsilon/d\tau > 0$ for $\tau < \tau_0$, and as $F_{rad} \propto d\varepsilon/d\tau$ for NS sources the radiative force always plays against gravity, unlike the case of BH sources.

Note that the spectra of NS sources both in the soft and hard state can be adequately fitted by single-temperature Comptonization models (Pazis et al. 2006, Falanga et al. 2005, Farinelli et al. 2008, Cocchi et al. 2010). This observational fact demonstrates that an assumption of isothermal plasma in the TL can be applicable to X-ray data analysis from NS binaries. The question is how one can estimate this average temperature of the TL which is, in fact, established by photon scattering and cooling processes.

In order to establish this average plasma temperature T_e one should estimate the mean energy density in the TL as

$$\langle \varepsilon(\tau) \rangle = \frac{1}{\tau_0} \int_0^{\tau_0} \varepsilon(\tau) d\tau = \frac{Q_{tot}}{c} (2 + \tau_0). \quad (8)$$

It is worth emphasizing the similarity between equations (3) and (8) in our paper and equation (13) in Bisnovatyi-Kogan et al. (1980) who studied the radiation emission due to gas accretion onto a NS. If we now substitute the result of equation (8) into equation (3), after a bit of straightforward algebra we obtain

$$\frac{kT_e\tau_0(2 + \tau_0)}{m_e c^2} = \frac{0.25}{1 + Q_{disk}/Q_{cor}}. \quad (9)$$

Keeping in mind the definition of the Compton parameter $Y \approx AN_{sc}$ (Rybicki & Lightman 1989), where $A \sim 4kT_e/m_e c^2$ and $N_{sc} \sim \text{Max}(\tau_0^2, \tau_0)$ are the average photon energy gain per scattering and average number of scatterings, respectively, we can rewrite equation (9) as follows

$$Y \sim \frac{1}{1 + Q_{disk}/Q_{cor}}. \quad (10)$$

Equation (10) is one of the main points of our theoretical model and shows that in the diffusion approximation the Compton parameter, which determines the spectral index, is just a function of the corona and disk cooling fluxes.

3.2. The radiative transfer model of small variations of index α during spectral state transition in NS sources

As we have shown in Section 2.1, the observed spectral index α of most NS LMXBs undergoes small variation around 1, namely $\alpha = 1 \pm 0.2$ when the electron temperature of Compton cloud varies from about 2.5 to 25 keV (see Fig. 1). Thus here

we propose a model the spectral formation in the TL (corona) which can explain the stability of index α if energy release in the disk is much less than that in the TL. Namely we show that $\alpha \approx 1 + O(Q_{\text{disk}}/Q_{\text{cor}})$.

As already pointed out in classical works (Sunyaev & Titarchuk 1980, 1985, hereafter ST85, Titarchuk 1994), spectral formation in plasma clouds of finite dimensions (bounded medium) is related to the distribution law of the number of scatterings that seed photons experience before escaping. If u_{av} denotes the average number of photon scatterings and the dimensionless scattering number is $u = N_e \sigma_T c t$, then the distribution law for $u \gg u_{av}$ is given by (see ST85)

$$P(u) = A(u_{av}, \tau_0) e^{-\beta u}. \quad (11)$$

For a diffusion regime when $\tau_0 \gtrsim 1.5$, it results $\beta = \lambda_1^2/3$, where λ_1 is the first eigenvalue of the diffusion space operator. As reported in ST85, the eigenvalue problem for photon diffusion in a slab with total optical depth $2\tau_0$ is derived from solution of the differential equation for the zero-moment intensity

$$\frac{d^2 J}{d\tau^2} + \lambda^2 J = 0, \quad (12)$$

with absorption boundary conditions $dJ/d\tau - (3/2)J = 0$ and $dJ/d\tau + (3/2)J = 0$, for $\tau = 0$ and $\tau = 2\tau_0$, respectively. This leads to the transcendental equation for the eigenvalue λ_n , $n = 1, 2, 3, \dots$

$$\tan(\lambda_n \tau_0) = \frac{2}{3\lambda_n}, \quad (13)$$

which has the solution for $\tau_0 \gg 1$ and $n = 1$

$$\lambda_1 = \frac{\pi}{2(\tau_0 + 2/3)}. \quad (14)$$

It is important to emphasize that the same result for λ_1 is obtained by solving equation (12) for a slab with total optical depth τ_0 but with reflection condition $dJ/d\tau = 0$ at $\tau = \tau_0$. This is not surprising as such a condition is actually met at the center of a symmetric slab with total optical depth $2\tau_0$ and $0 \leq \tau \leq 2\tau_0$. Thus, the same mathematical result is obtained for two different geometrical configurations. In the first case (symmetric slab with total optical depth $2\tau_0$) it represents, e.g., an accretion disk (ST85 treatment), in our present case we are dealing with a boundary layer with total optical depth τ_0 , whose asymmetry is due to the presence of a reflector (NS surface) at one of the two boundaries. In both cases one obtains

$$\beta = \frac{\pi^2}{12(\tau_0 + 2/3)^2}. \quad (15)$$

Generalizing to the case of arbitrary optical depth τ_0 , the diffusion operator $L_{\text{diff}} = (1/3)d^2 J/d\tau^2$ is replaced by the radiative transfer operator L_τ applied to $J(\tau)$ (see ST85 and TL95) which for the disk geometry is

$$L_\tau J = \frac{1}{2} \int_0^{2\tau_0} J(\tau') E_1(|\tau - \tau'|) d\tau' - J, \quad (16)$$

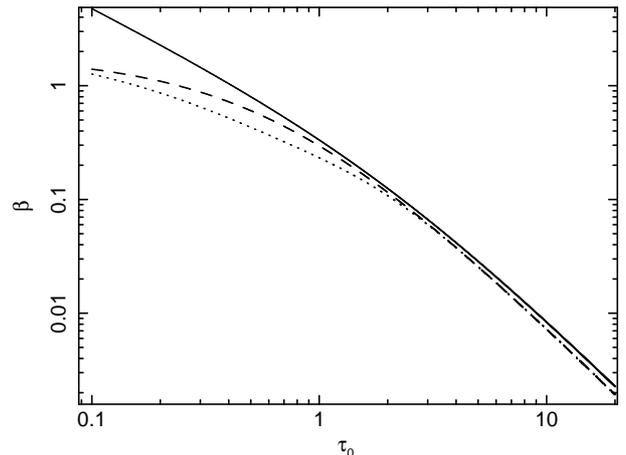


Fig. 2. Values of the β -parameter as a function of the optical depth τ_0 . The solid, dashed and dotted lines correspond to definition of β given in equations (15), (17) and (18), respectively.

where $E_1(z)$ is the exponential integral of the first order. In this case, the derived value for β is (Titarchuk 1994, TL95)

$$\beta = \frac{\pi^2}{12(\tau_0 + 2/3)^2} (1 - e^{-1.35\tau_0}) + 0.45e^{-3.7\tau_0} \ln \frac{10}{3\tau_0}. \quad (17)$$

Now having in mind equation (9), we introduce the parameter

$$\beta_{\text{diff}} = \frac{1}{\tau_0(2 + \tau_0)}, \quad (18)$$

and in Figure 2 we show the values of β for the cases reported in formulas (15), (17) and (18) as a function of optical depth τ_0 . It is possible to see that actually for $\tau_0 \gtrsim 1.5$ all values of β are practically close each other, but they deviate for $\tau_0 \lesssim 1$. For example their difference is about 30% for $\tau_0=1$.

Using the definition of α (see Eq. 2), where β is replaced by β_{diff} (Eq. 18), and equation (9), we obtain the diffusion spectral index as

$$\alpha_{\text{diff}} = -\frac{3}{2} + \sqrt{\frac{9}{4} + \frac{1 + Q_{\text{disk}}/Q_{\text{cor}}}{0.25}}, \quad (19)$$

or $\alpha_{\text{diff}} = 1 + 0.8 Q_{\text{disk}}/Q_{\text{cor}}$ for $Q_{\text{disk}}/Q_{\text{cor}} < 1$.

Thus as it follows from Eq. (19), in the diffusion regime the TC spectral index can be expressed in terms of $Q_{\text{disk}}/Q_{\text{cor}}$ (the intercepted disk over coronal fluxes), instead of TL electron temperature kT_e and optical depth τ_0 (see Eqs. [2] and [15]). In Figure 3 we present a plot of α_{diff} as a function of $Q_{\text{disk}}/Q_{\text{cor}}$, which shows that it ranges from 1 to 1.6 as $Q_{\text{disk}}/Q_{\text{cor}}$ increases from 0 to 1. One can see the observable values of index $\alpha \sim 1$ takes place if the energy release in the disk is much less than in TL, namely if $Q_{\text{disk}}/Q_{\text{cor}} \ll 1$.

4. Results and discussion

In this Paper we compare observational data of a sample of NS sources (Fig. 1) with the theoretical results which follows from the radiative transfer model in the diffusion approximation. The data show that *independently of the source spectral state*, which we have parametrized through the measured TL

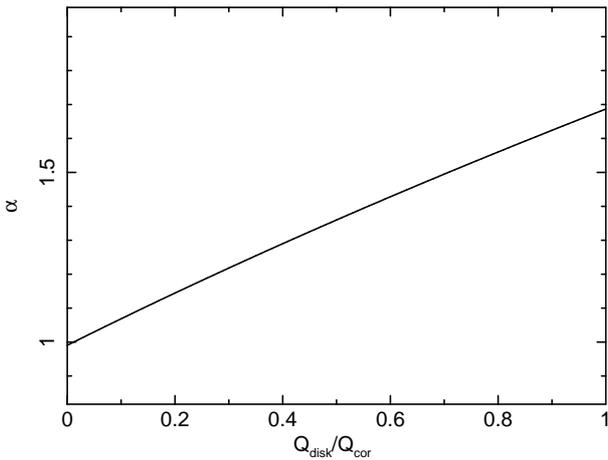


Fig. 3. Theoretical thermal Comptonization index α as a function of the ratio $Q_{\text{disk}}/Q_{\text{cor}}$ according to equation (19).

electron temperature kT_e , the spectral index $\alpha = 1 \pm 0.2$. We derived an estimate of the energy index α for TC spectra in NS LMXBs using an equation for the diffusion approximation in a slab geometry with the reflection (100 %) boundary condition, which is valid for optical depth $\tau_0 \gtrsim 1.5$. In particular, we find that in this approximation it is possible to express the value of α as a function of the ratio of the flux from the accretion disk intercepted by the corona (TL, see Eq. [19]) and the energy release in the corona itself. The agreement of the model with the data takes place when the condition $Q_{\text{disk}}/Q_{\text{cor}} \ll 1$ holds (see Fig. 3).

This behavior of spectral index α actually has important consequences on putting constraints in the accretion geometry of NS LMXBs. In fact, as already pointed-out in the Introduction, the broad-band persistent X-ray spectra of these sources are usually fitted by a two-component model consisting of BB-like emission plus a strong TC bump. Both the cases for which the BB temperature (kT_{bb}) is lower or higher than that of the seed photons of the TC bump (kT_s) provide equally acceptable fits. In the first case ($kT_{\text{bb}} < kT_s$), the origin of the direct BB component is attributed to the accretion disk, in the second case ($kT_{\text{bb}} > kT_s$) to the NS surface. This dichotomy in the spectral analysis has not been overcome for a long time. Some help in this direction has come with the discovery of the transient X-ray tails in some of the brightest sources (see references in the Introduction): in fact, if the origin of the hard PL-like X-ray emission is attributed to a combined thermal plus bulk (converging flow) effect in the innermost part of the TL close to the NS surface, then it becomes natural to suggest that the observed direct BB photon spectrum mostly originates in the TL/NS surface region, providing the seed photons for bulk Comptonization (see Fig. 1 in Farinelli et al. 2007).

The theoretical results derived in this Paper, in our opinion strengthen this scenario. In fact, when computing the total energetic spectral budget of the sources in the 0.1-40 keV (where most of the emission is produced) the dominating TC bump carries out more than 70% of the source luminosity, the remaining part due to the direct BB component. If this BB originates close to the NS surface, then it is evident that the disk contribution to

the X-ray luminosity is very small. Given that Q_{disk} in equation (19) represents the flux from the accretion disk intercepted by the corona and thus is smaller or at most equal to the directly emitted part, this eventually leads to $\alpha \sim 1$. In this framework, we can make some considerations about the higher value of α (~ 1.6) measured for GX 354-0 (see Fig. 1). Di Salvo et al. (2000b) fitted the broad-band spectrum of the source with a BB plus COMPTT model (Titarchuk 1994) with $kT_{\text{bb}} < kT_s$, a modelization corresponding to the case in which a significantly higher fraction of the X-ray luminosity comes from the accretion disk. This would turn of course in an enhanced value of Q_{disk} , and looking at Figure 3 it is straightforward to see that increasing $Q_{\text{disk}}/Q_{\text{cor}}$ corresponds to increasing α . Actually, it would be interesting to see what it does happen by fitting the *BeppoSAX* spectra of GX 354-0 with the same model but with $kT_{\text{bb}} > kT_s$. Note also in Fig. 1 that for 1E 1724-3045 and two spectral states of GS 1826-238, $\alpha \sim 1$ with electron temperature $kT_e \gtrsim 20$ keV. Using equation (9) and the best-fit values of kT_e reported for the two sources, with $Q_{\text{disk}}/Q_{\text{cor}} \sim 0$, we obtain $\tau_0 \sim 1.3$ for 1E 1724-3045 and $\tau_0 \sim 1.6-1.7$ for GS 1826-238, respectively. These values of the optical depth allow actually to deal with the diffusion approximation within degree of accuracy still satisfactory as can be seen from Fig. 2. Moreover, it is worth emphasizing that for a slab with optical depth τ_0 and inner reflection boundary condition, photons which are back-scattered from the reflecting surface before escaping experience in fact an optical depth $\sim 2\tau_0$, which further enhances the diffusion approximation validity.

The other issue to point out is that the lower limit on α (for the extreme case $Q_{\text{disk}}/Q_{\text{cor}} = 0$) derived by our model is 1, while there is a handful of sources for which $0.8 < \alpha < 1$. Different reasons can in fact concur to this result. First of all, it is well known that multi-component modeling of X-ray spectra may have some influence in the determination of the best-fit parameters. In particular, in the energy band where TC dominates ($\lesssim 30$ keV) Gaussian emission lines around 6.4-6.7 keV are often observed and inclusion of narrow-feature component in the model can affect the continuum parameters. Additional biases can come from the energy-band coverage (in particular when using *RXTE* data, which start from about 3 keV) and uncertainties in the calibration of the instrumental effective area, which may play some role in particular for spectra which are far away from being powerlaw-like.

In terms of theoretical predictions, our analytical model is intended to provide a description of the observed stability of the spectral index but may of course have some limitations, in particular the energy-independent treatment of the radiation field and the simple slab approximation for the TL geometry. However it is possible that the vertical height-scale H of the TL is higher than its radial extension, it is also likely that there is some dependence of H on the radial distance from the NS surface. Moreover, in the pure plane-parallel geometry photons are allowed to escape only through the surface of the slab, while in this case the slab has limited extension and photons presumably can escape also from its “walls”.

5. Conclusions

We reported results on the value of the thermal Comptonization (TC) spectral index α for a sample of LMXBs NS sources in different spectral cases and we found that a part from GX 354–0 it lies in a belt around 1 ± 0.2 . We proposed a simple theoretical model using the diffusion approximation where α is found to be a function only of the ratio of the disk and corona fluxes. In particular when $Q_{\text{disk}}/Q_{\text{cor}} \ll 1$ it results $\alpha \sim 1$, which is consistent with observations. We are hopeful that our work will encourage to significantly extend the sample of observed NS sources using archival data and observations from present and future missions, in particular using as broad as possible energy band, especially below 3 keV in order to avoid biases in the spectral results. We also claim that our model can be helpful in solving the dichotomy related to the fact that equally good fit are obtained for the cases in which the observed direct BB component has a temperature higher or lower than that of the seed photons subjected to TC.

Acknowledgements. The authors are grateful to the referee whose suggestions strongly improved the quality of the paper with respect to the first version. This work was supported by grant from Italian PRIN-INAF 2007, “Bulk motion Comptonization models in X-ray Binaries: from phenomenology to physics”, PI M. Cocchi.

References

- Barret, D., Olive, J. F., Boirin, L., et al. 2000, *ApJ*, 533, 329
- Bisnovatyi-Kogan, G. S., Khlopov, M. Y., Chechetkin, V. M., & Eramzhyan, R. A. 1980, *SvA*, vol. 24, p. 716,
- Cocchi, M., Farinelli, R., Paizis, A., & Titarchuk, L. 2010, *A&A*, 509, 2
- D’Amico, F., Heindl, W. A., Rothschild, R. E., et al. 2001, *ApJ*, 547, L147
- Di Salvo, T., Goldoni, P., Stella, L., et al. 2006, *ApJ*, 649, L91
- Di Salvo, T., Farinelli, R., Burderi, L., et al. 2002, *A&A*, 386, 535
- Di Salvo, T., Robba, N. R., Iaria, R., et al. 2001, *ApJ*, 554, 49
- Di Salvo, T.; Stella, L.; Robba, N. R., et al. 2000a, 554, L119
- Di Salvo, T., Iaria, R., Burderi, L., & Robba, N.R. 2000b, *ApJ*, 542, 1034
- Falanga, M., Godtz, D., Goldoni, P., Farinelli, R., Goldwurm, A., et al. 2006, *A&A*, 458, 21
- Farinelli, R., Paizis, A., Landi, R., & Titarchuk, L. 2009, *A&A*, 498, 509
- Farinelli, R., Titarchuk, L., Paizis, A., & Frontera, F. 2008, *ApJ*, 680, 602, (F08)
- Farinelli, R., Titarchuk, L., & Frontera, F. 2007, *ApJ*, 662, 1167
- Gierliński, M., & Done, C. 2002, *MNRAS*, 337, 1373
- Haardt, F., & Maraschi, L. 1993, *ApJ*, 413, 50
- Hua, X-M & Titarchuk, L. 1995, *ApJ*, 449, 188
- Laurent, P. & Titarchuk, L. 1999, *ApJ*, 511, 289
- Lavagetto, G., Iaria, R., di Salvo, T., et al. 2004, *Nucl. Phys. B, Proc. Suppl.*, 132, 616
- Mitsuda, K., Inoue, H., Koyama, K., et al. 1984, *PASJ*, 36, 741
- Mitsuda, K., Inoue, H., Nakamura, N., & Tanaka, Y. 1989, *PASJ*, 41, 97
- Montanari, E., Titarchuk, L., & Frontera, F. 2009, *ApJ*, 692, 1597
- Oosterbroek, T., Parmar, A.N., Sidoli, L., in’t Zand, J.J.M., & Heise, J. 2001, *A&A*, 370, 532
- Paizis, A., Farinelli, R., Titarchuk, L., et al. 2006, *A&A*, 459, 187
- Paizis, A., Ebisawa, K., Tikkanen, T., et al. 2005, *A&A*, 435, 599
- Popham, R. & Sunyaev, R. A. 2001, *ApJ*, 547, 355
- Rybicki, G. B., & Lightman, A. P. 1979, *Radiative processes in Astrophysics* (New York: Wiley)
- Sunyaev, R. A. & Titarchuk, L. 1980, *A&A*, 86, 121
- Sunyaev, R. A. & Titarchuk, L. 1985, *A&A*, 143, 374
- Shaposhnikov, N., & Titarchuk, L. 2009, *ApJ*, 699, 453
- Titarchuk, L., & Seifina, E., 2009, *ApJ*, 706, 1463
- Titarchuk, L., & Shaposhnikov, N. 2005, *ApJ*, 626, 298
- Titarchuk, L., & Fiorito, R. 2004, *ApJ*, 612, 988
- Titarchuk, L., & Osherovich, V. 1999, *ApJ*, 518, L95
- Titarchuk, L., Lapidus, J., & Muslimov, A. 1998, *ApJ*, 499, 315
- Titarchuk, L., Mastichiadis, A., & Kylafis, N. D. 1997, *ApJ*, 487, 831
- Titarchuk, L., & Lyubarskij, Y. 1995, *ApJ*, 450, 876 (TL95)
- Titarchuk, L. 1994, *ApJ*, 434, 570
- White, N. E., Peacock, A., Hasinger, G., et al. 1986, *MNRAS*, 218, 129
- White, N. E., Stella, L., & Parmar, A. 1988, *ApJ*, 324, 363
- Zel’dovich, Ya.B. & Shakura, N.I. 1969, *Sov. Astronomy* 13, 175